

# Results and perspectives of DAMA/LIBRA

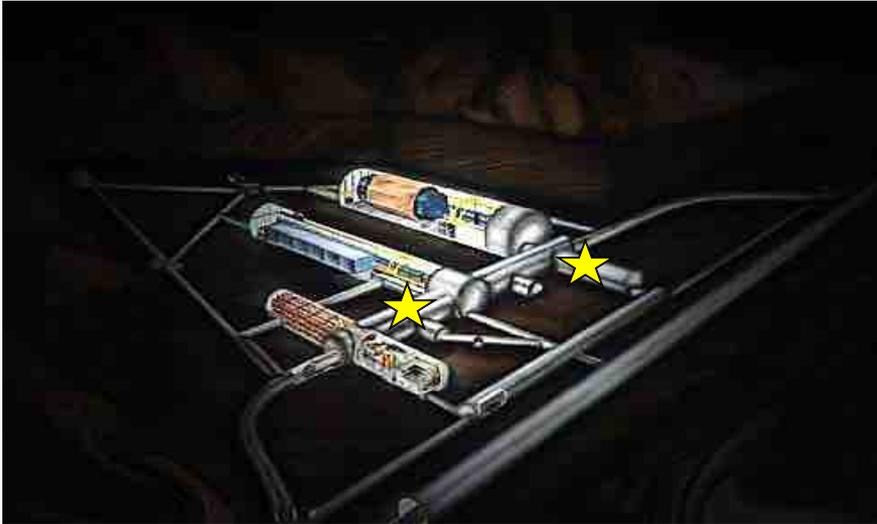


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MEDEX'17  
Prague, CZ  
May 29 – June 2, 2017

# DAMA set-ups

an observatory for rare processes @ LNGS



- DAMA/LIBRA (DAMA/NaI)
- DAMA/LXe
- DAMA/R&D
- DAMA/Crys
- DAMA/Ge

For other results  
see the talks of  
O. Polishchuk,  
R. Cerulli,

## Collaboration:

Roma Tor Vergata, Roma La Sapienza, LNGS, IHEP/Beijing  
+ by-products and small scale expts.: INR-Kiev + other institutions  
+ neutron meas.: ENEA-Frascati, ENEA-Casaccia  
+ in some studies on  $\beta\beta$  decays (DST-MAE and Inter-Universities project):  
IIT Kharagpur and Ropar, India

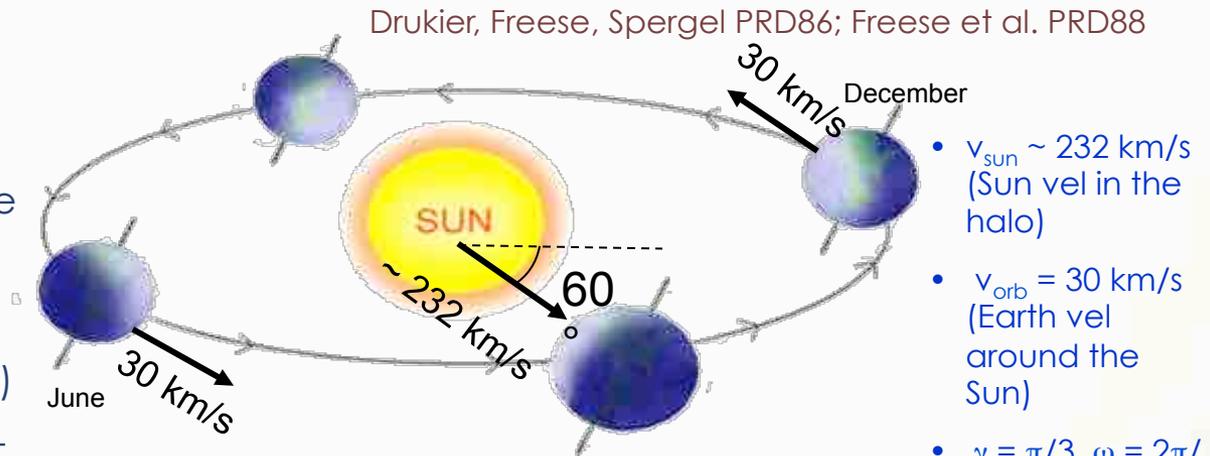
web site: <http://people.roma2.infn.it/dama>

# The annual modulation: a model independent signature for the investigation of DM particles component in the galactic halo

With the present technology, the annual modulation is the main model independent signature for the DM signal. Although the modulation effect is expected to be relatively small a suitable large-mass, low-radioactive set-up with an efficient control of the running conditions can point out its presence.

## Requirements:

- 1) Modulated rate according cosine
- 2) In low energy range
- 3) With a proper period (1 year)
- 4) With proper phase (about 2 June)
- 5) Just for single hit events in a multi-detector set-up
- 6) With modulation amplitude in the region of maximal sensitivity must be <7% for usually adopted halo distributions, but it can be larger in case of some possible scenarios



$$v_{\oplus}(t) = v_{\text{sun}} + v_{\text{orb}} \cos\gamma \cos[\omega(t-t_0)]$$

$$S_k[\eta(t)] = \int_{\Delta E_k} \frac{dR}{dE_R} dE_R \cong S_{0,k} + S_{m,k} \cos[\omega(t-t_0)]$$

the DM annual modulation signature has a different origin and peculiarities (e.g. the phase) than those effects correlated with the seasons

To mimic this signature, spurious effects and side reactions must not only - obviously - be able to account for the whole observed modulation amplitude, but also to satisfy contemporaneously all the requirements

# The pioneer DAMA/NaI: ≈100 kg highly radiopure NaI(Tl)

## Performances:

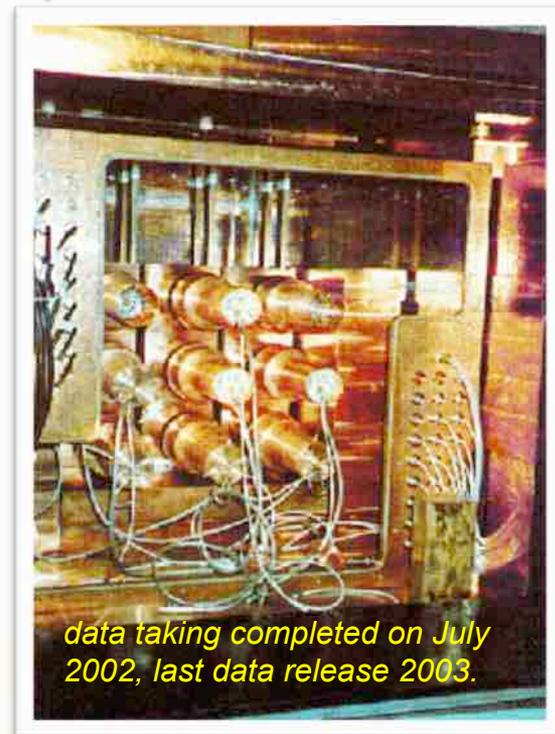
N.Cim.A112(1999)545-575, EPJC18(2000)283,  
Riv.N.Cim.26 n. 1(2003)1-73, IJMPD13(2004)2127

## Results on rare processes:

- Possible Pauli exclusion principle violation PLB408(1997)439
- CNC processes PRC60(1999)065501
- Electron stability and non-paulian transitions in Iodine atoms (by L-shell) PLB460(1999)235
- Search for solar axions PLB515(2001)6
- Exotic Matter search EPJdirect C14(2002)1
- Search for superdense nuclear matter EPJA23(2005)7
- Search for heavy clusters decays EPJA24(2005)51

## Results on DM particles:

- PSD PLB389(1996)757
- Investigation on diurnal effect N.Cim.A112(1999)1541
- Exotic Dark Matter search PRL83(1999)4918
- **Annual Modulation Signature** PLB424(1998)195, PLB450(1999)448, PRD61(1999)023512, PLB480(2000)23, EPJC18(2000)283, PLB509(2001)197, EPJC23(2002)61, PRD66(2002)043503, Riv.N.Cim.26 n.1 (2003)1, IJMPD13(2004)2127, IJMPA21(2006)1445, EPJC47(2006)263, IJMPA22(2007)3155, EPJC53(2008)205, PRD77(2008)023506, MPLA23(2008)2125



**Model independent evidence of a particle DM  
component in the galactic halo at  $6.3\sigma$  C.L.**

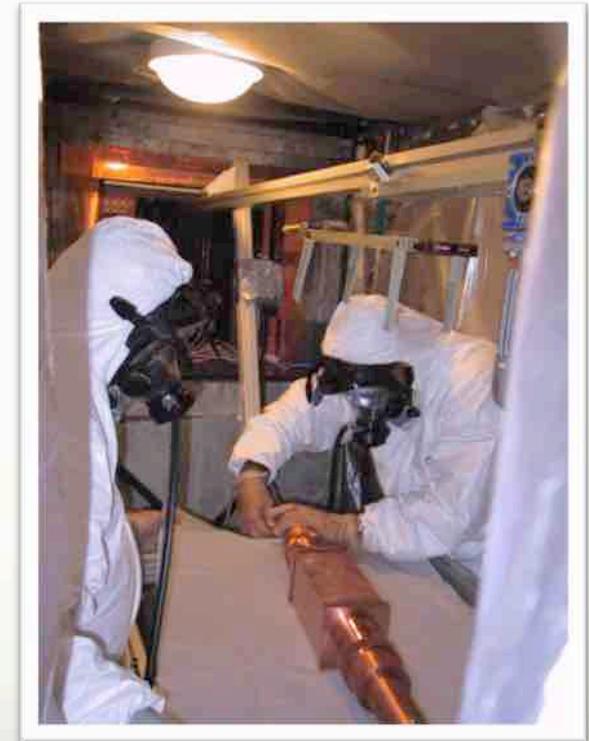
total exposure (7 annual cycles) 0.29 ton×yr

# The DAMA/LIBRA set-up ~250 kg NaI(Tl) (Large sodium Iodide Bulk for RARE processes)

As a result of a 2nd generation R&D for more radiopure NaI(Tl) by exploiting new chemical/physical radiopurification techniques (all operations involving - including photos - in HP Nitrogen atmosphere)



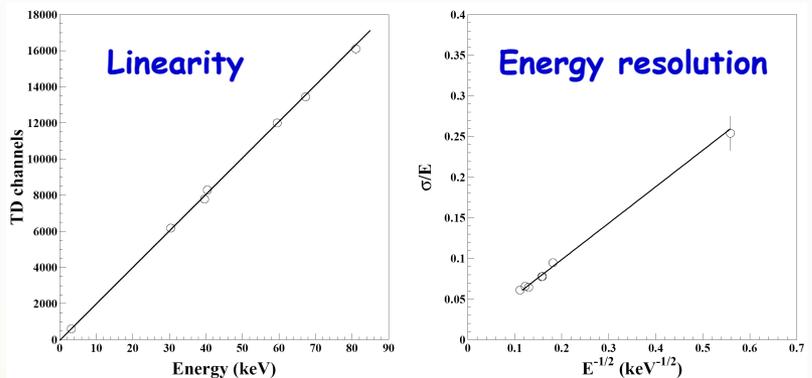
Residual contaminations in the new DAMA/LIBRA NaI(Tl) detectors:  $^{232}\text{Th}$ ,  $^{238}\text{U}$  and  $^{40}\text{K}$  at level of  $10^{-12}$  g/g



- Radiopurity, performances, procedures, etc.: NIMA592(2008)297, JINST 7 (2012) 03009
- Results on DM particles, **Annual Modulation Signature**: EPJC56(2008)333, EPJC67(2010)39, EPJC73(2013)2648.  
**Related results**: PRD84(2011)055014, EPJC72(2012)2064, IJMPA28(2013)1330022, EPJC74(2014)2827, EPJC74(2014)3196, EPJC75(2015)239, EPJC75(2015)400, IJMPA31(2016) dedicated issue, EPJC77(2017)83
- Results on rare processes: **PEPv**: EPJC62(2009)327; **CNC**: EPJC72(2012)1920; **IPP in  $^{241}\text{Am}$** : EPJA49(2013)64

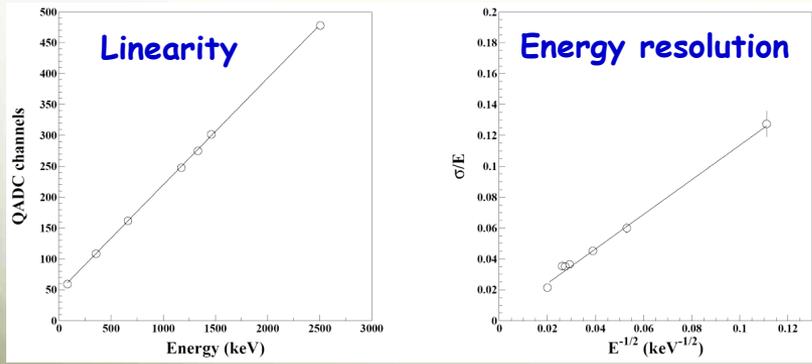
# DAMA/LIBRA calibrations

Low energy: various external gamma sources ( $^{241}\text{Am}$ ,  $^{133}\text{Ba}$ ) and internal X-rays or gamma's ( $^{40}\text{K}$ ,  $^{125}\text{I}$ ,  $^{129}\text{I}$ ), routine calibrations with  $^{241}\text{Am}$



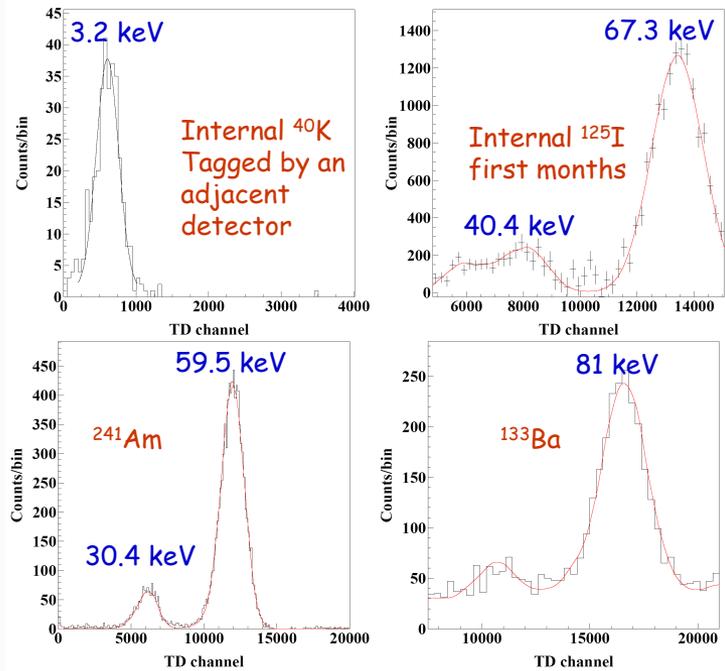
$$\frac{\sigma_{LE}}{E} = \frac{(0.448 \pm 0.035)}{\sqrt{E(\text{keV})}} + (9.1 \pm 5.1) \cdot 10^{-3}$$

High energy: external sources of gamma rays (e.g.  $^{137}\text{Cs}$ ,  $^{60}\text{Co}$  and  $^{133}\text{Ba}$ ) and gamma rays of 1461 keV due to  $^{40}\text{K}$  decays in an adjacent detector, tagged by the 3.2 keV X-rays

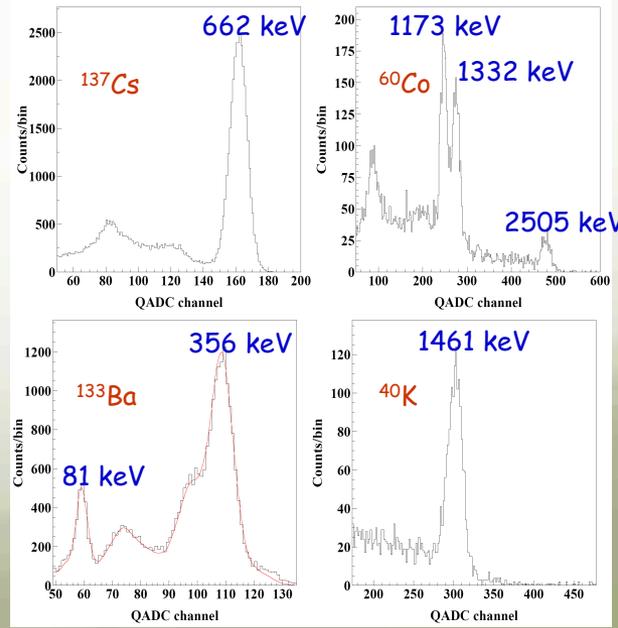


$$\frac{\sigma_{HE}}{E} = \frac{(1.12 \pm 0.06)}{\sqrt{E(\text{keV})}} + (17 \pm 23) \cdot 10^{-4}$$

Thus, here and hereafter keV means keV electron equivalent



The curves superimposed to the experimental data have been obtained by simulations



# Complete DAMA/LIBRA-phase1

	Period	Mass (kg)	Exposure (kg×day)	$(\alpha - \beta^2)$
DAMA/LIBRA-1	Sept. 9, 2003 - July 21, 2004	232.8	51405	0.562
DAMA/LIBRA-2	July 21, 2004 - Oct. 28, 2005	232.8	52597	0.467
DAMA/LIBRA-3	Oct. 28, 2005 - July 18, 2006	232.8	39445	0.591
DAMA/LIBRA-4	July 19, 2006 - July 17, 2007	232.8	49377	0.541
DAMA/LIBRA-5	July 17, 2007 - Aug. 29, 2008	232.8	66105	0.468
DAMA/LIBRA-6	Nov. 12, 2008 - Sept. 1, 2009	242.5	58768	0.519
DAMA/LIBRA-7	Sept. 1, 2009 - Sept. 8, 2010	242.5	62098	0.515
DAMA/LIBRA-phase1	Sept. 9, 2003 - Sept. 8, 2010		379795	1.04 ton×yr
DAMA/NaI + DAMA/LIBRA-phase1:				1.33 ton×yr

a ton × yr experiment? done

- EPJC56(2008)333
- EPJC67(2010)39
- EPJC73(2013)2648
- calibrations: ≈96 Mevents from sources
- acceptance window eff: 95 Mevents (≈3.5 Mevents/keV)

## DAMA/LIBRA-phase1:

- First upgrade on Sept 2008: replacement of some PMTs in HP N<sub>2</sub> atmosphere, new Digitizers (U1063A Acqiris 1GS/s 8-bit High-speed cPCI), new DAQ system with optical read-out installed

## DAMA/LIBRA-phase2 (running):

- Second upgrade at end 2010: replacement of all the PMTs with higher Q.E. ones from dedicated developments
- commissioning on 2011
  - Goal: lowering the software energy threshold
- Fall 2012: new preamplifiers installed + special trigger modules. Other new components in the electronic chain in development

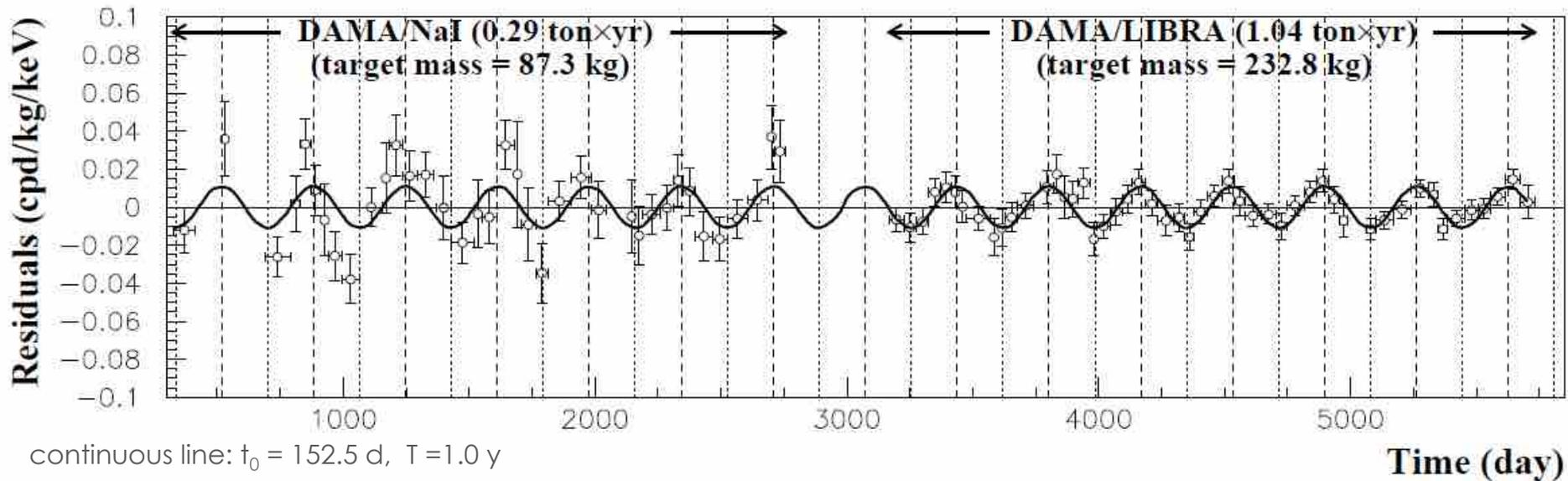


# Model Independent Annual Modulation Result

DAMA/NaI + DAMA/LIBRA-phase1 Total exposure: 487526 kg×day = **1.33 ton×yr**

EPJC 56(2008)333,  
EPJC 67(2010)39,  
EPJC 73(2013)2648

## Single-hit residuals rate vs time in 2-6 keV



Absence of modulation? No

$$\chi^2/\text{dof} = 154/87$$

$$P(A=0) = 1.3 \times 10^{-5}$$

$$\text{Fit: } t_0 = 152.5 \text{ d, } T = 1.0 \text{ y}$$

$$A = (0.0110 \pm 0.0012) \text{ cpd/kg/keV}$$

$$\chi^2/\text{dof} = 70.4/86 \quad 9.2 \sigma \text{ C.L.}$$

The data favor the presence of a modulated behaviour with all the proper features for DM particles in the galactic halo at about  $9.2\sigma$  C.L.

# Model Independent Annual Modulation Result

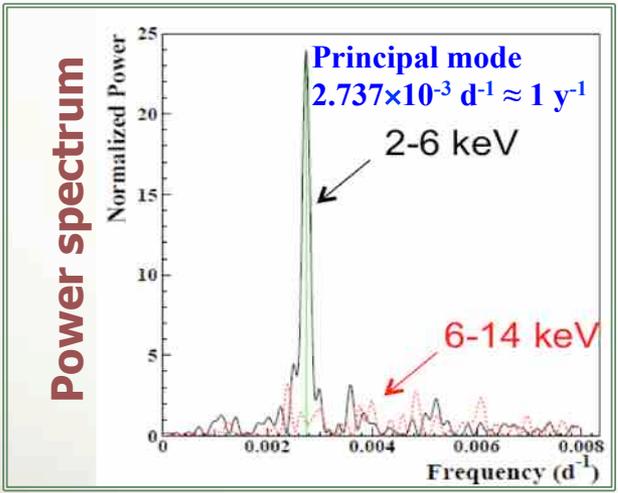
**DAMA/NaI + DAMA/LIBRA-phase1** Total exposure: 487526 kg×day = **1.33 ton×yr**

EPJC 56(2008)333, EPJC 67(2010)39, EPJC 73(2013)2648

The measured modulation amplitudes (A), period (T) and phase ( $t_0$ ) from the single-hit residual rate vs time

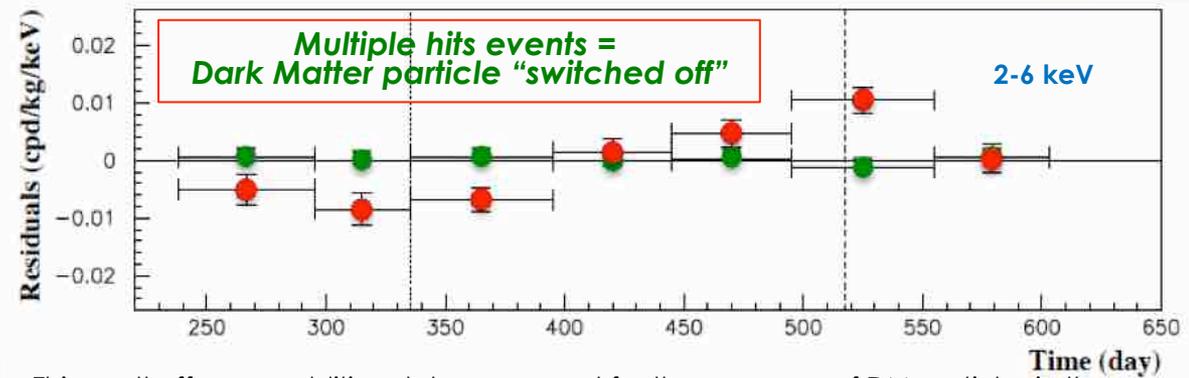
	A(cpd/kg/keV)	T=2 $\pi$ / $\omega$ (yr)	$t_0$ (day)	C.L.
<b>DAMA/NaI+DAMA/LIBRA-phase1</b>				
(2-4) keV	<b>0.0190 ±0.0020</b>	<b>0.996 ±0.002</b>	<b>134 ± 6</b>	<b>9.5<math>\sigma</math></b>
(2-5) keV	<b>0.0140 ±0.0015</b>	<b>0.996 ±0.002</b>	<b>140 ± 6</b>	<b>9.3<math>\sigma</math></b>
(2-6) keV	<b>0.0112 ±0.0012</b>	<b>0.998 ±0.002</b>	<b>144 ± 7</b>	<b>9.3<math>\sigma</math></b>

$\text{Acos}[\omega(t-t_0)]$



No systematics or side reaction able to account for the measured modulation amplitude and to satisfy all the peculiarities of the signature

Comparison between **single hit residual rate (red points)** and **multiple hit residual rate (green points)**; Clear modulation in the single hit events; No modulation in the residual rate of the multiple hit events  
**A = -(0.0005 ± 0.0004) cpd/kg/keV**



This result offers an additional strong support for the presence of DM particles in the galactic halo further excluding any side effect either from hardware or from software procedures or from background

The data favor the presence of a modulated behaviour with all the proper features for DM particles in the galactic halo at about 9.2 $\sigma$  C.L.

# Model Independent Annual Modulation Result

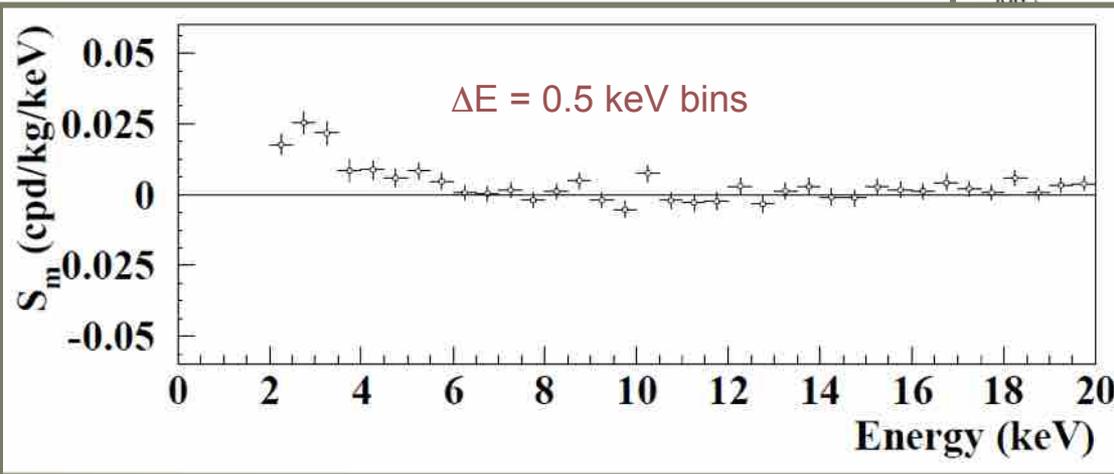
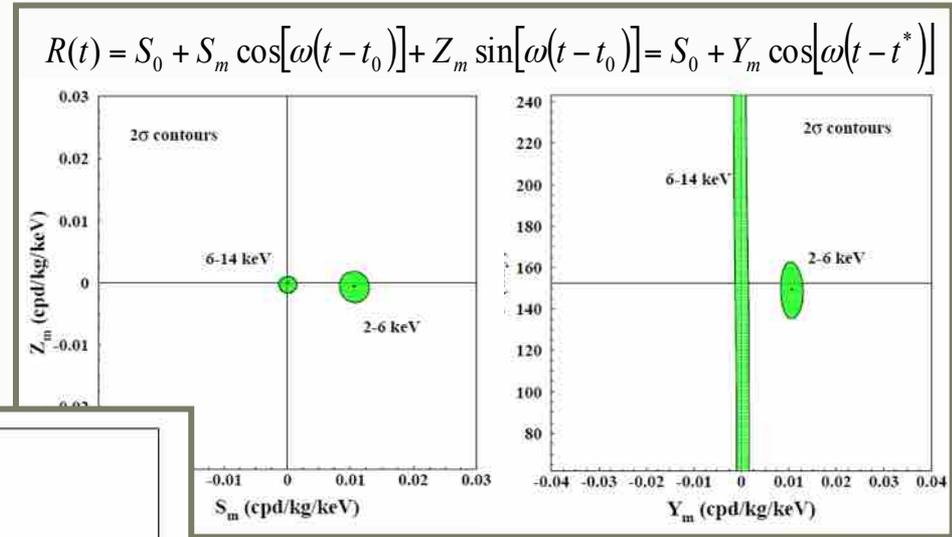
**DAMA/NaI + DAMA/LIBRA-phase1** Total exposure: 487526 kg×day = **1.33 ton×yr**

EPJC 56(2008)333, EPJC 67(2010)39, EPJC 73(2013)2648

- No modulation above 6 keV
- No modulation in the whole energy spectrum
- No modulation in the 2-6 keV multiple-hit events

$$R(t) = S_0 + S_m \cos[\omega(t - t_0)]$$

here  $T = 2\pi/\omega = 1$  yr and  $t_0 = 152.5$  day



A clear modulation is present in the (2-6) keV energy interval, while  $S_m$  values compatible with zero are present just above

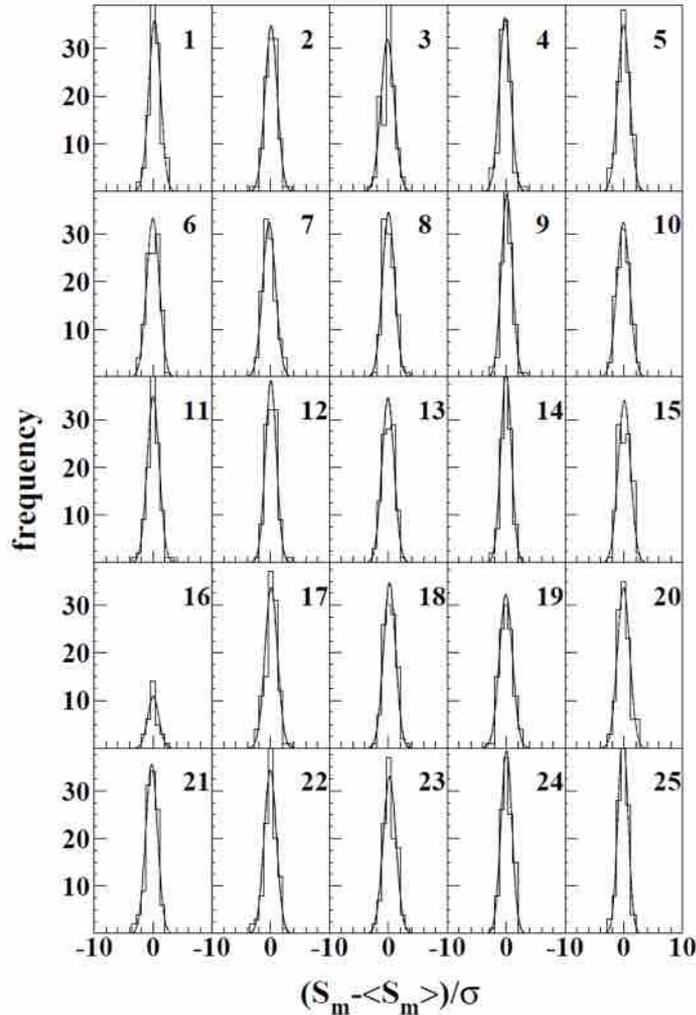
The  $S_m$  values in the (6-20) keV energy interval have random fluctuations around zero with  $\chi^2$  equal to 35.8 for 28 degrees of freedom (upper tail probability 15%)

# Statistical distributions of the modulation amplitudes ( $S_m$ )

- a)  $S_m$  for each detector, each annual cycle and each considered energy bin (here 0.25 keV)
- b)  $\langle S_m \rangle$  = mean values over the detectors and the annual cycles for each energy bin;  $\sigma$  = error on  $S_m$

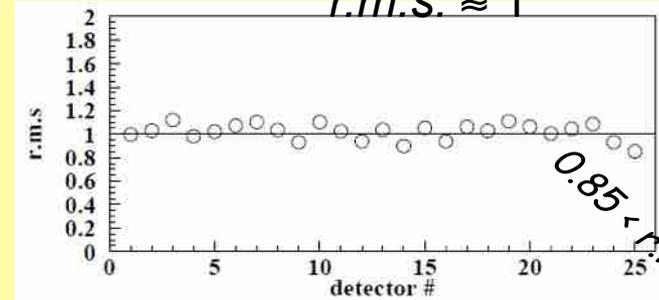
**DAMA/LIBRA-phase1 (7 years)**  
**total exposure: 1.04 ton $\times$ yr**

Each panel refers to each detector separately; 112 entries = 16 energy bins in 2-6 keV energy interval  $\times$  7 DAMA/LIBRA-phase1 annual cycles (for crys 16, 2 annual cycle, 32 entries)



2-6 keV

Standard deviations of  $(S_m - \langle S_m \rangle) / \sigma$  for each detectors  
*r.m.s.*  $\approx 1$



$$x = (S_m - \langle S_m \rangle) / \sigma,$$

$$\chi^2 = \sum x^2$$

Individual  $S_m$  values follow a normal distribution since  $(S_m - \langle S_m \rangle) / \sigma$  is distributed as a Gaussian with a unitary standard deviation (r.m.s.)

$\Rightarrow$   $S_m$  statistically well distributed in all the detectors, energy bin and annual cycles

# Summary of the results obtained in the additional investigations of possible systematics or side reactions – DAMA/LIBRA-phase1

(NIMA592(2008)297, EPJC56(2008)333, J. Phys. Conf. ser. 203(2010)012040, arXiv:0912.0660, S.I.F. Atti Conf.103(211), Can. J. Phys. 89 (2011) 11, Phys.Proc.37(2012)1095, EPJC72(2012)2064, arxiv:1210.6199 & 1211.6346, IJMPA28(2013)1330022, EPJC74(2014)3196)

Source	Main comment	Cautious upper limit (90%C.L.)
<b>RADON</b>	Sealed Cu box in HP Nitrogen atmosphere, 3-level of sealing, etc.	$<2.5 \times 10^{-6}$ cpd/kg/keV
<b>TEMPERATURE</b>	Installation is air conditioned+ detectors in Cu housings directly in contact with multi-ton shield → huge heat capacity + T continuously recorded	$<10^{-4}$ cpd/kg/keV
<b>NOISE</b>	Effective full noise rejection near threshold	$<10^{-4}$ cpd/kg/keV
<b>ENERGY SCALE</b>	Routine + intrinsic calibrations	$<1-2 \times 10^{-4}$ cpd/kg/keV
<b>EFFICIENCIES</b>	Regularly measured by dedicated calibrations	$<10^{-4}$ cpd/kg/keV
<b>BACKGROUND</b>	No modulation above 6 keV; no modulation in the (2-6) keV <i>multiple-hits</i> events; this limit includes all possible sources of background	$<10^{-4}$ cpd/kg/keV
<b>SIDE REACTIONS</b>	Muon flux variation measured at LNGS	$<3 \times 10^{-5}$ cpd/kg/keV

+ they cannot satisfy all the requirements of annual modulation signature

Thus, they cannot mimic the observed annual modulation effect

# No role for $\mu$ in DAMA annual modulation result

## ✓ Direct $\mu$ interaction in DAMA/LIBRA set-up:

DAMA/LIBRA surface  $\approx 0.13 \text{ m}^2$   
 $\mu$  flux @ DAMA/LIBRA  $\approx 2.5 \mu/\text{day}$

It cannot mimic the signature: already excluded by  $R_{90}$ , by *multi-hits* analysis + different phase, etc.

## ✓ Rate, $R_n$ , of fast neutrons produced by $\mu$ :

- $\Phi_\mu$  @ LNGS  $\approx 20 \mu \text{ m}^{-2}\text{d}^{-1}$  ( $\pm 1.5\%$  modulated)
- Annual modulation amplitude at low energy due to  $\mu$  modulation:

$$S_m(\mu) = R_n g \varepsilon f_{\Delta E} f_{\text{single}} 2\% / (M_{\text{setup}} \Delta E)$$

Moreover, this modulation also induces a variation in other parts of the energy spectrum and in the *multi-hits* events

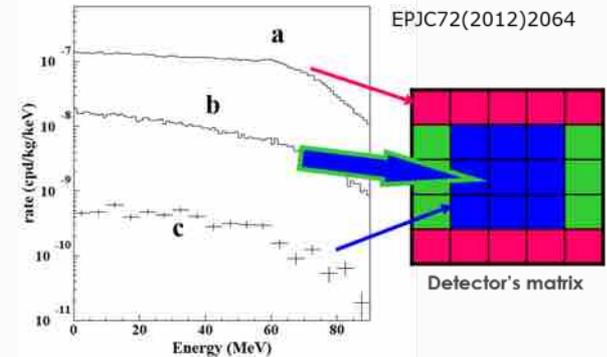
## ✓ Inconsistency of the phase between DAMA signal and $\mu$ modulation

$\mu$  flux @ LNGS (MACRO, LVD, BOREXINO)  $\approx 3 \cdot 10^{-4} \text{ m}^{-2}\text{s}^{-1}$ ;  
 modulation amplitude 1.5%; **phase: July  $7 \pm 6$  d, June  $29 \pm 6$  d** (Borexino)

The DAMA phase: **May  $26 \pm 7$  days** (stable over 13 years)

The DAMA phase is  $5.7\sigma$  far from the LVD/BOREXINO phases of muons ( $7.1\sigma$  far from MACRO measured phase)

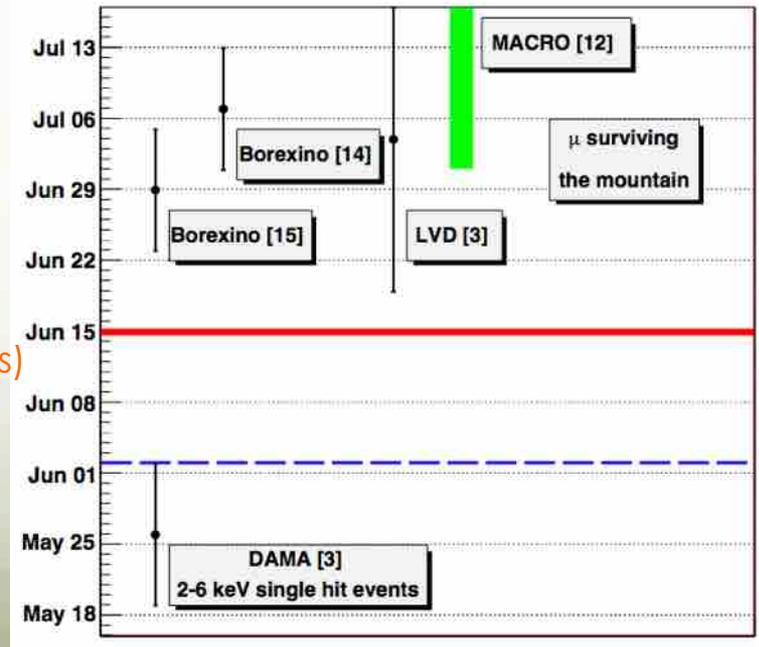
... many others arguments EPJC72(2012)2064, EPJC74(2014)3196



MonteCarlo simulation

$$S_m(\mu) < (0.3-2.4) \times 10^{-5} \text{ cpd/kg/keV}$$

It cannot mimic the signature: already excluded by  $R_{90}$ , by *multi-hits* analysis + different phase, etc.



- Contributions to the total **neutron flux** at LNGS;
- **Counting rate** in DAMA/LIBRA for *single-hit* events, in the (2 - 6) keV energy region induced by:

- neutrons,
- muons,
- solar neutrinos.

$$\Phi_k = \Phi_{0,k} (1 + \eta_k \cos\omega (t - t_k))$$

$$R_k = R_{0,k} (1 + \eta_k \cos\omega (t - t_k))$$

EPJC 74 (2014) 3196 (also EPJC 56 (2008) 333, EPJC 72 (2012) 2064, IJMPA 28 (2013) 1330022)

Modulation amplitudes

Source	$\Phi_{0,k}^{(n)}$ (neutrons cm <sup>-2</sup> s <sup>-1</sup> )	$\eta_k$	$t_k$	$R_{0,k}$ (cpd/kg/keV)	$A_k = R_{0,k}\eta_k$ (cpd/kg/keV)	$A_k/S_m^{exp}$	
SLOW neutrons	thermal n (10 <sup>-2</sup> - 10 <sup>-1</sup> eV)	1.08 × 10 <sup>-6</sup> [15]	≈ 0 however ≪ 0.1 [2, 7, 8]	-	< 8 × 10 <sup>-6</sup> [2, 7, 8]	≪ 8 × 10 <sup>-7</sup>	≪ 7 × 10 <sup>-5</sup>
	epithermal n (eV-keV)	2 × 10 <sup>-6</sup> [15]	≈ 0 however ≪ 0.1 [2, 7, 8]	-	< 3 × 10 <sup>-3</sup> [2, 7, 8]	≪ 3 × 10 <sup>-4</sup>	≪ 0.03
FAST neutrons	fission, (α, n) → n (1-10 MeV)	≈ 0.9 × 10 <sup>-7</sup> [17]	≈ 0 however ≪ 0.1 [2, 7, 8]	-	< 6 × 10 <sup>-4</sup> [2, 7, 8]	≪ 6 × 10 <sup>-5</sup>	≪ 5 × 10 <sup>-3</sup>
	μ → n from rock (> 10 MeV)	≈ 3 × 10 <sup>-9</sup> (see text and ref. [12])	0.0129 [23]	end of June [23, 7, 8]	≪ 7 × 10 <sup>-4</sup> (see text and [2, 7, 8])	≪ 9 × 10 <sup>-6</sup>	≪ 8 × 10 <sup>-4</sup>
	μ → n from Pb shield (> 10 MeV)	≈ 6 × 10 <sup>-9</sup> (see footnote 3)	0.0129 [23]	end of June [23, 7, 8]	≪ 1.4 × 10 <sup>-3</sup> (see text and footnote 3)	≪ 2 × 10 <sup>-5</sup>	≪ 1.6 × 10 <sup>-3</sup>
	ν → n (few MeV)	≈ 3 × 10 <sup>-10</sup> (see text)	0.03342 *	Jan. 4th *	≪ 7 × 10 <sup>-5</sup> (see text)	≪ 2 × 10 <sup>-6</sup>	≪ 2 × 10 <sup>-4</sup>
direct μ	$\Phi_0^{(\mu)} \approx 20 \mu \text{ m}^{-2} \text{ d}^{-1}$ [20]	0.0129 [23]	end of June [23, 7, 8]	≈ 10 <sup>-7</sup> [2, 7, 8]	≈ 10 <sup>-9</sup>	≈ 10 <sup>-7</sup>	
direct ν	$\Phi_0^{(\nu)} \approx 6 \times 10^{10} \nu \text{ cm}^{-2} \text{ s}^{-1}$ [26]	0.03342 *	Jan. 4th *	≈ 10 <sup>-5</sup> [31]	3 × 10 <sup>-7</sup>	3 × 10 <sup>-5</sup>	

\* The annual modulation of solar neutrino is due to the different Sun-Earth distance along the year; so the relative modulation amplitude is twice the eccentricity of the Earth orbit and the phase is given by the perihelion.

All are negligible w.r.t. the annual modulation amplitude observed by DAMA/LIBRA and they cannot contribute to the observed modulation amplitude.

+ In no case neutrons (of whatever origin) can mimic the DM annual modulation signature since some of the peculiar requirements of the signature would fail, such as the neutrons would induce e.g. variations in all the energy spectrum, variation in the multiple hit events,... which were not observed.

# Model-independent evidence by DAMA/NaI and DAMA/LIBRA

well compatible with several candidates in many astrophysical, nuclear and particle physics scenarios

Neutralino as LSP in various SUSY theories

Various kinds of WIMP candidates with several different kind of interactions  
Pure SI, pure SD, mixed + Migdal effect + channeling,... (from low to high mass)

a heavy  $\nu$  of the 4-th family

Pseudoscalar, scalar or mixed light bosons with axion-like interactions

WIMP with preferred inelastic scattering

Mirror Dark Matter

Light Dark Matter

Dark Matter (including some scenarios for WIMP) electron-interacting

Sterile neutrino

Self interacting Dark Matter

heavy exotic candidates, as "4th family atoms", ...

Elementary Black holes such as the Daemons



Kaluza Klein particles

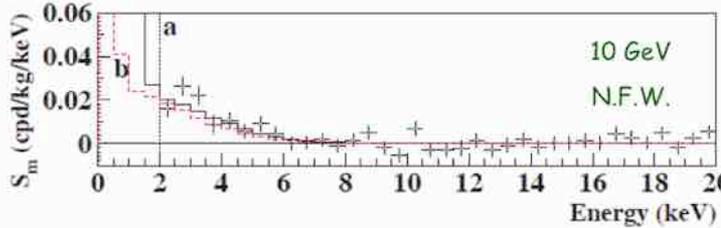
... and more

# Model-independent evidence by DAMA/NaI and DAMA/LIBRA

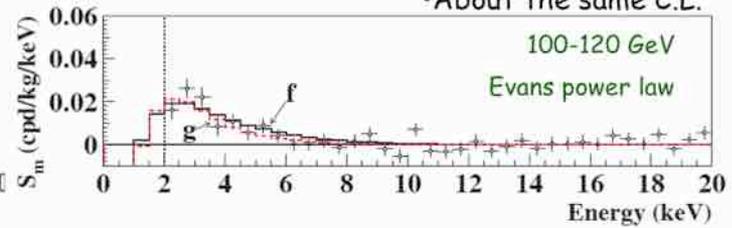
well compatible with several candidates in many astrophysical, nuclear and particle physics scenarios

Just few examples of interpretation of the annual modulation in terms of candidate particles in some scenarios

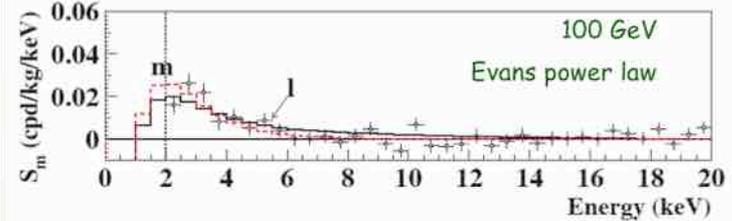
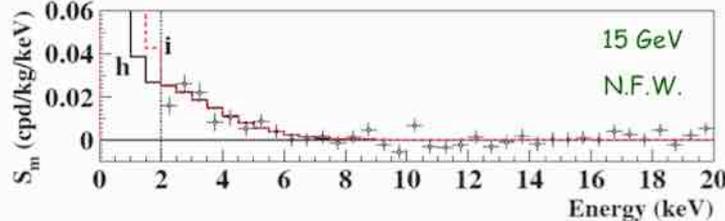
WIMP: SI



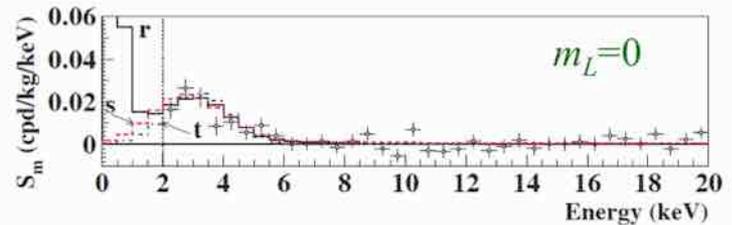
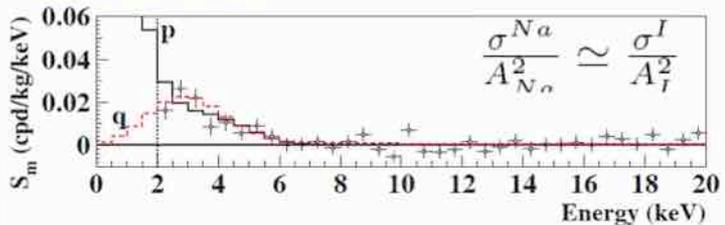
•Not best fit  
•About the same C.L.



WIMP: SI & SD  $\theta = 2.435$



LDM, bosonic DM

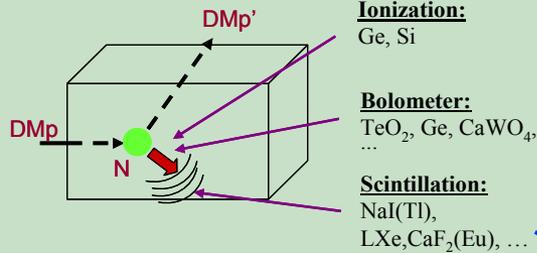


Compatibility with several candidates;  
other ones are open

# Some direct detection processes:

- Scatterings on nuclei

→ detection of nuclear recoil energy



- Inelastic Dark Matter:  $W + N \rightarrow W^* + N$

→ W has 2 mass states  $\chi^+$ ,  $\chi^-$  with  $\delta$  mass splitting

→ Kinematical constraint for the inelastic scattering of  $\chi^-$  on a nucleus

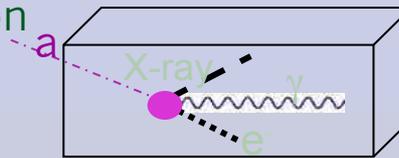
$$\frac{1}{2} \mu v^2 \geq \delta \Leftrightarrow v \geq v_{thr} = \sqrt{\frac{2\delta}{\mu}}$$

- Excitation of bound electrons in scatterings on nuclei

→ detection of recoil nuclei + e.m. radiation

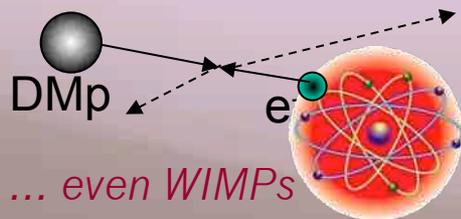
- Conversion of particle into e.m. radiation

→ detection of  $\gamma$ , X-rays,  $e^-$



- Interaction only on atomic electrons

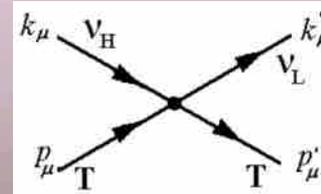
→ detection of e.m. radiation



- Interaction of light DMP (LDM) on  $e^-$  or nucleus with production of a lighter particle

→ detection of electron/nucleus recoil energy

e.g. sterile  $\nu$



e.g. signals from these candidates are **completely lost** in experiments based on “rejection procedures” of the e.m. component of their rate

... also other ideas ...



# About interpretations and comparisons

See e.g.: Riv.N.Cim.26 n.1(2003)1, IJMPD13(2004)2127, EPJC47(2006)263, IJMPA21(2006)1445, EPJC56(2008)333, PRD84(2011)055014, IJMPA28(2013)1330022

## ...and experimental aspects...

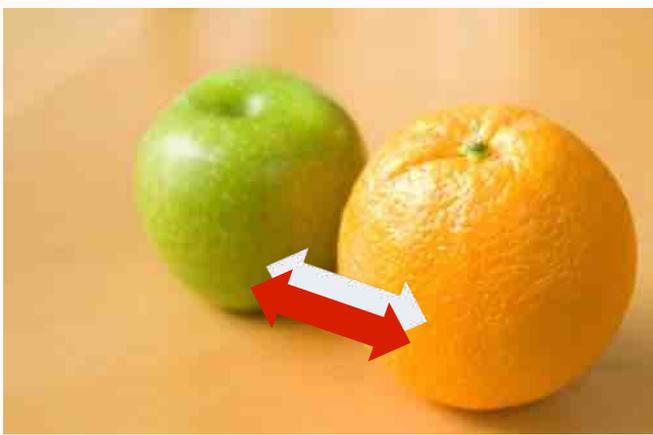
- Exposures
- Energy threshold
- Detector response (phe/keV)
- Energy scale and energy resolution
- Calibrations
- Stability of all the operating conditions.
- Selections of detectors and of data.
- Subtraction/rejection procedures and stability in time of all the selected windows and related quantities
- Efficiencies
- Definition of fiducial volume and non-uniformity
- Quenching factors, channeling, ...
- ...

## ...models...

- Which particle?
- Which interaction coupling?
- Which Form Factors for each target-material?
- Which Spin Factor?
- Which nuclear model framework?
- Which scaling law?
- Which halo model, profile and related parameters?
- Streams?
- ...

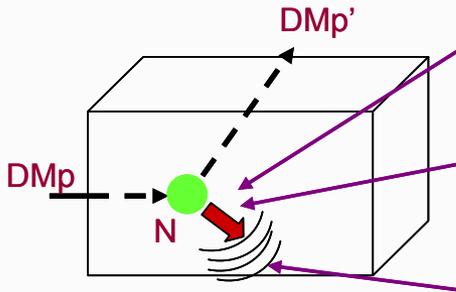
Uncertainty in experimental parameters, as well as necessary assumptions on various related astrophysical, nuclear and particle-physics aspects, affect all the results at various extent, both in terms of exclusion plots and in terms of allowed regions/volumes. Thus comparisons with a fixed set of assumptions and parameters' values are intrinsically strongly uncertain.

**No experiment can be directly compared in model independent way with DAMA**



# ... an example in literature...

Case of DM particles inducing elastic scatterings on target-nuclei, SI case

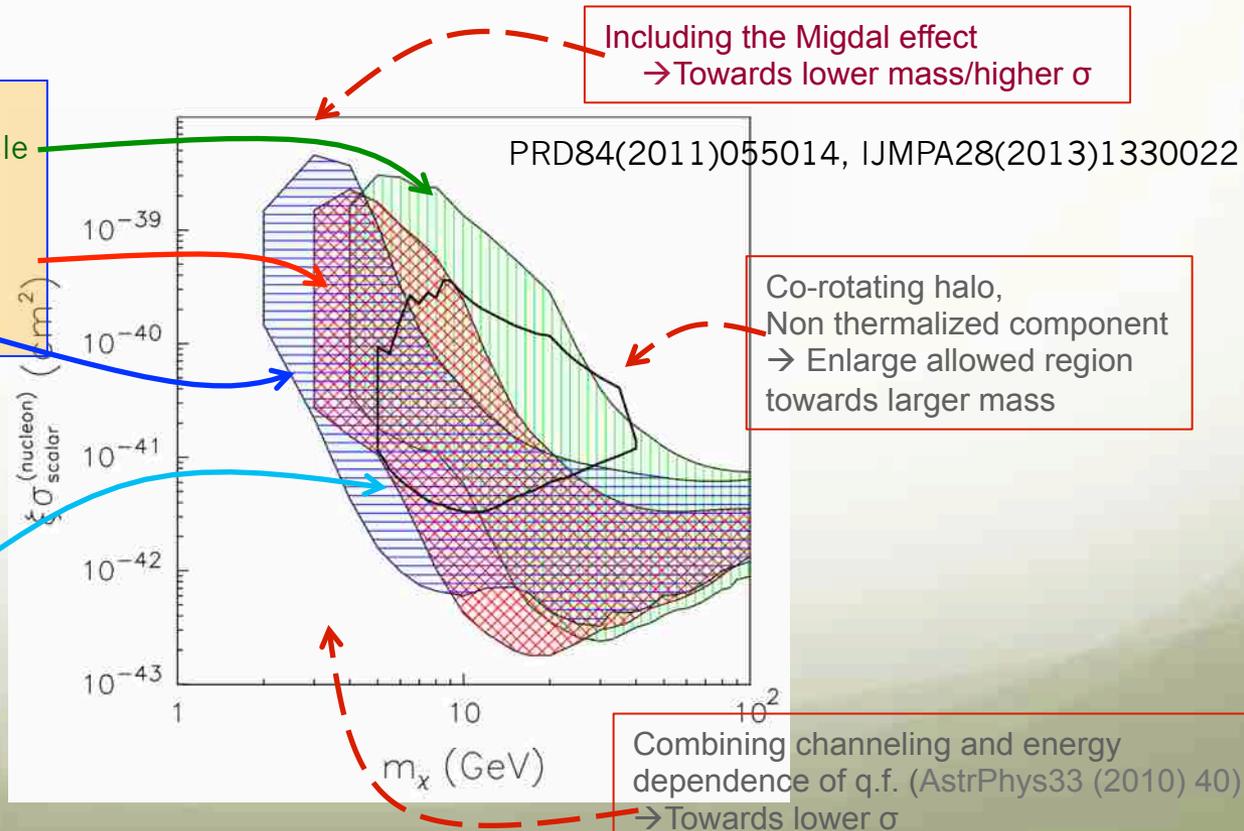


## Regions in the nucleon cross section vs DM particle mass plane

- Some velocity distributions and uncertainties considered.
- The DAMA regions represent the domain where the likelihood-function values differ more than  $7.5\sigma$  from the null hypothesis (absence of modulation).
- For CoGeNT a fixed value for the Ge quenching factor and a Helm form factor with fixed parameters are assumed.
- The CoGeNT region includes configurations whose likelihood-function values differ more than  $1.64\sigma$  from the null hypothesis (absence of modulation). This corresponds roughly to 90% C.L. far from zero signal.

DAMA allowed regions for a particular set of astrophysical, nuclear and particle Physics assumptions without (green), with (blue) channeling, with energy-dependent Quenching Factors (red);  $7.5 \sigma$  C.L.

CoGeNT; qf at fixed assumed value  $1.64 \sigma$  C.L.



## Scratching Below the Surface of the Most General Parameter Space (S. Scopel arXiv:1505.01926)

Most general approach: consider ALL possible NR couplings, including those depending on velocity and momentum

- A much wider parameter space opens up

- First explorations show that indeed large rooms for compatibility can be achieved

$$\begin{aligned} \mathcal{O}_1 &= 1_{\chi} 1_N, \\ \mathcal{O}_2 &= (v^\perp)^2, \\ \mathcal{O}_3 &= i \vec{S}_N \cdot \left( \frac{\vec{q}}{m_N} \times \vec{v}^\perp \right), \\ \mathcal{O}_4 &= \vec{S}_\chi \cdot \vec{S}_N, \\ \mathcal{O}_5 &= i \vec{S}_\chi \cdot \left( \frac{\vec{q}}{m_N} \times \vec{v}^\perp \right), \\ \mathcal{O}_6 &= \left( \vec{S}_\chi \cdot \frac{\vec{q}}{m_N} \right) \left( \vec{S}_N \cdot \frac{\vec{q}}{m_N} \right), \\ \mathcal{O}_7 &= \vec{S}_N \cdot \vec{v}^\perp, \\ \mathcal{O}_8 &= \vec{S}_\chi \cdot \vec{v}^\perp, \\ \mathcal{O}_9 &= i \vec{S}_\chi \cdot \left( \vec{S}_N \times \frac{\vec{q}}{m_N} \right), \\ \mathcal{O}_{10} &= i \vec{S}_N \cdot \frac{\vec{q}}{m_N}, \\ \mathcal{O}_{11} &= i \vec{S}_\chi \cdot \frac{\vec{q}}{m_N}. \end{aligned}$$

... and much more considering experimental and theoretical uncertainties

## Other examples

DMp with preferred inelastic interaction:  
 $\chi^- + N \rightarrow \chi^+ + N$

- iDM mass states  $\chi^+$ ,  $\chi^-$  with  $\delta$  mass splitting
- Kinematic constraint for iDM:

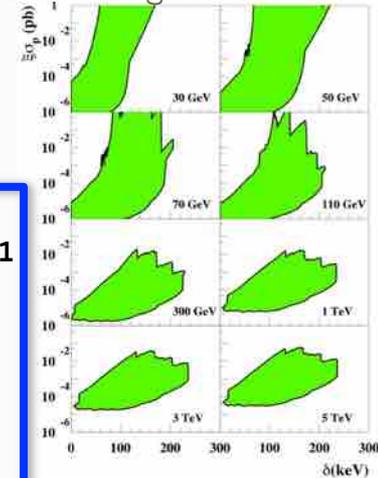
$$\frac{1}{2} \mu v^2 \geq \delta \Leftrightarrow v \geq v_{thr} = \sqrt{\frac{2\delta}{\mu}}$$

iDM interaction on TI nuclei of the NaI(Tl) dopant?

PRL106(2011)011301

- For large splittings, the dominant scattering in NaI(Tl) can occur off of Thallium nuclei, with  $A \sim 205$ , which are present as a dopant at the  $10^{-3}$  level in NaI(Tl) crystals.
- large splittings do not give rise to sizeable contribution on Na, I, Ge, Xe, Ca, O, ... nuclei.

DAMA/NaI+DAMA/LIBRA  
 Slices from the 3d allowed volume in given scenario



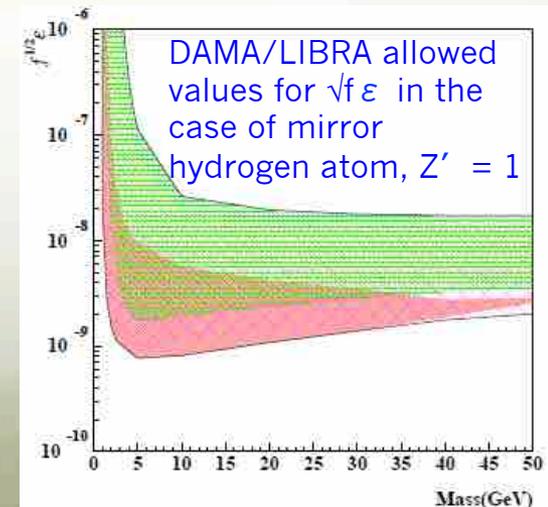
Fund. Phys. 40(2010)900

## Mirror Dark Matter

Asymmetric mirror matter: mirror parity spontaneously broken  $\Rightarrow$  mirror sector becomes a heavier and deformed copy of ordinary sector  
 (See EPJC75(2015)400)

- Interaction portal: photon - mirror photon kinetic mixing  $\frac{\epsilon}{2} F^{\mu\nu} F'_{\mu\nu}$
- mirror atom scattering of the ordinary target nuclei in the NaI(Tl) detectors of DAMA/LIBRA set-up with the Rutherford-like cross sections.

$$\sqrt{f} \cdot \epsilon \quad \text{coupling const. and fraction of mirror atom}$$



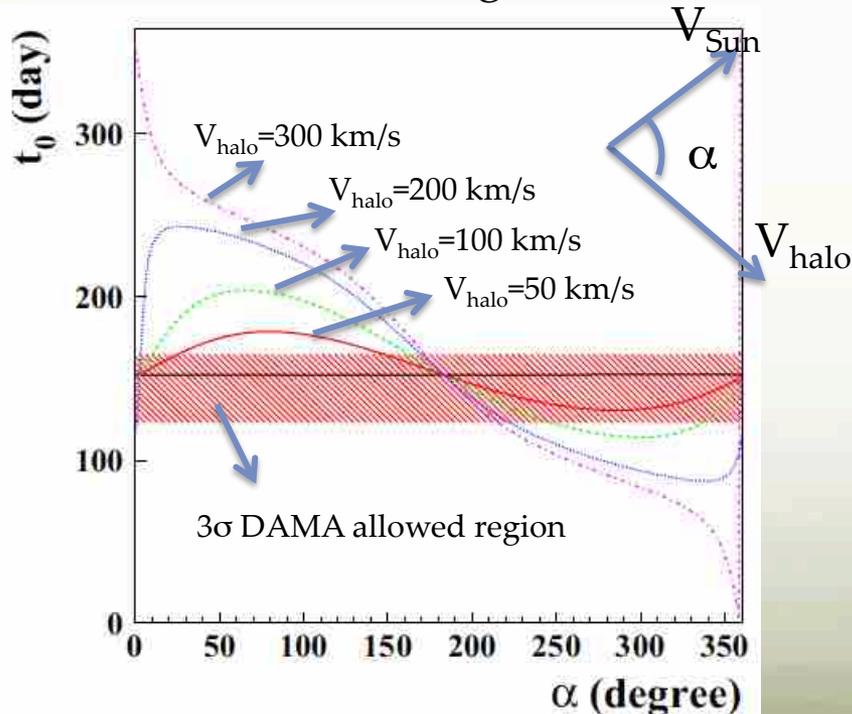
# DAMA annual modulation effect and Symmetric mirror matter

Eur. Phys. J. C (2017) 77

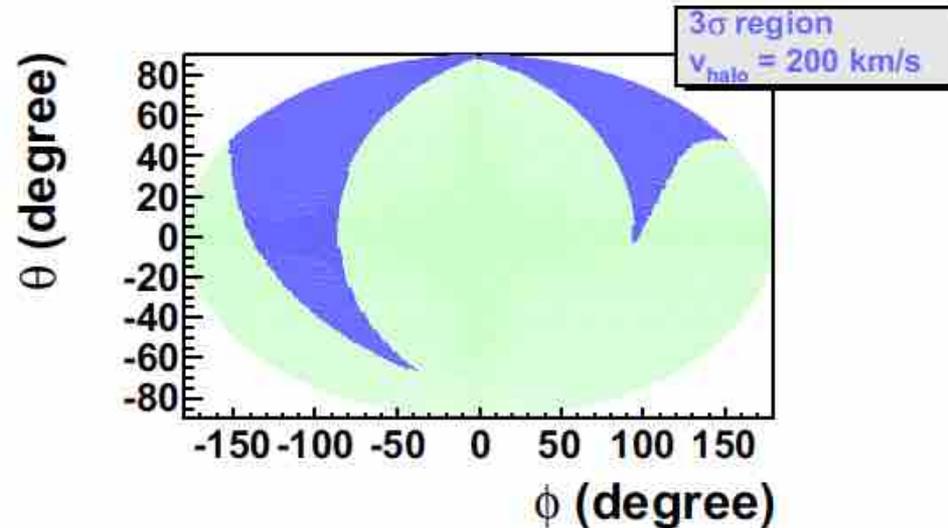
Symmetric mirror matter:

- halo composed by a bubble of Mirror particles of different species; Sun is travelling across the bubble which is moving in the Galactic Frame (GF);
- the mirror particles in the bubble have Maxwellian velocity distribution in a frame where the bubble is at rest; cold and hot bubble with temp from  $10^4$  K to  $10^8$  K)
- interaction via photon - mirror photon kinetic mixing

Examples of expected phase of the annual modulation signal



The blue regions correspond to directions of the halo velocities in GC ( $\theta, \phi$ ) giving a phase compatible at  $3\sigma$  with DAMA phase



# DAMA annual modulation effect and Symmetric mirror matter

Eur. Phys. J. C (2017) 77

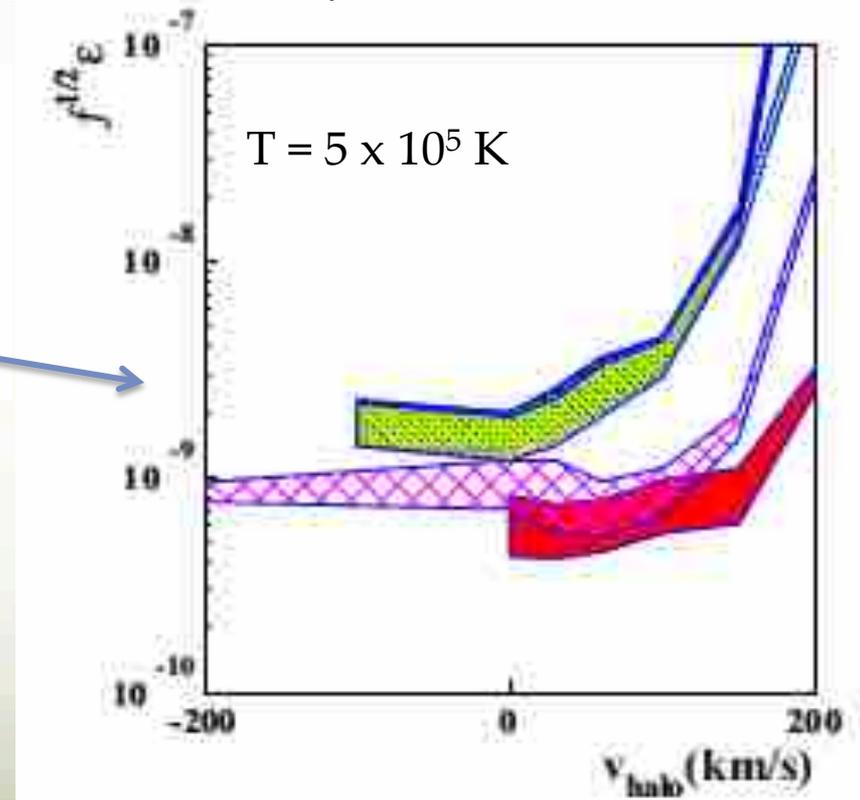
Symmetric mirror matter:

- Results refers to halo velocities parallel or anti-parallel to the Sun ( $\alpha = 0, \pi$ ). For these configurations the expected phase is June 2
- The only parameter whose value will be varied in the analysis is the  $V_{\text{halo}}$  module (positive velocity will correspond to halo moving in the same direction of the Sun while negative velocity will correspond to opposite direction)

Mirror matter composition	H (%)	He (%)	C (%)	O (%)	Fe (%)
H', He'	25	75	-	-	-
H', He', C', O'	12.5	75	7	5.5	-
H', He', C', O', Fe'	20	74	0.9	5	0.1

DAMA/LIBRA allowed values for  $\sqrt{f \epsilon}$  in different scenarios

$\sqrt{f \cdot \epsilon}$  coupling const. and DM fraction as mirror atom



Many configurations and halo models favoured by the DAMA annual modulation effect corresponds to couplings values well compatible with cosmological bounds.

# Perspectives for the future

## *Other signatures?*

- *Diurnal effects*
- *Second order effects*
- *Shadow effects*
- *Directionality*
- *...*

# Diurnal effects in DAMA/LIBRA-phase1

EPJC 74 (2014) 2827

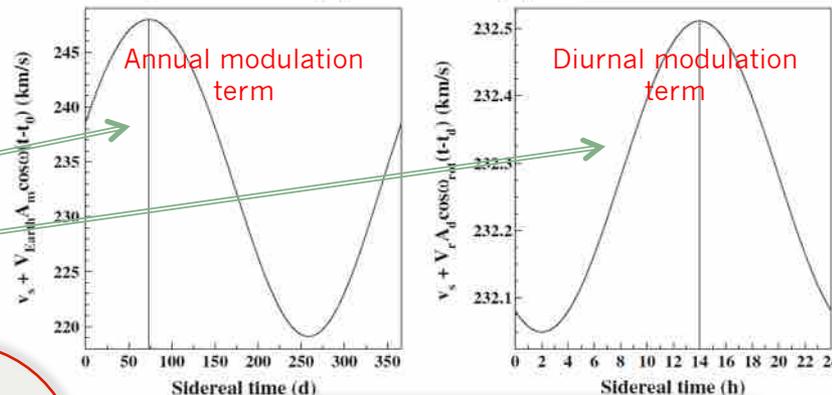
A diurnal effect with the sidereal time is expected for DM because of Earth rotation

Velocity of the detector in the terrestrial laboratory:  $\vec{v}_{lab}(t) = \vec{v}_{LSR} + \vec{v}_{\odot} + \vec{v}_{rev}(t) + \vec{v}_{rot}(t)$ ,

Since:

- $|\vec{v}_s| = |\vec{v}_{LSR} + \vec{v}_{\odot}| \approx 232 \pm 50$  km/s,
- $|\vec{v}_{rev}(t)| \approx 30$  km/s
- $|\vec{v}_{rot}(t)| \approx 0.34$  km/s at LNGS

$$v_{lab}(t) \simeq v_s + \hat{v}_s \cdot \vec{v}_{rev}(t) + \hat{v}_s \cdot \vec{v}_{rot}(t).$$



Expected signal counting rate in a given k-th energy bin:

$$S_k[v_{lab}(t)] \simeq S_k[v_s] + \left[ \frac{\partial S_k}{\partial v_{lab}} \right]_{v_s} [V_{Earth} B_m \cos \omega(t - t_0) + V_r B_d \cos \omega_{rot}(t - t_d)]$$

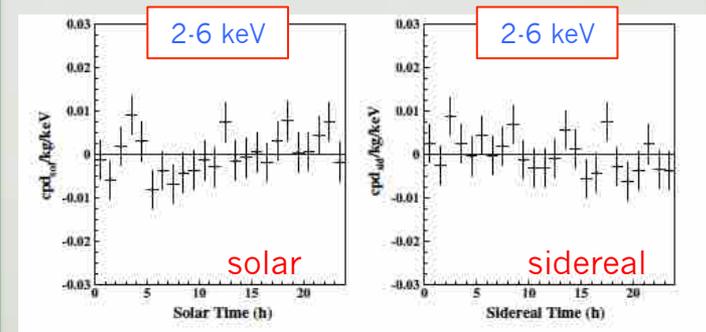
The ratio  $R_{dy}$  is a model independent constant:

$$R_{dy} = \frac{S_d}{S_m} = \frac{V_r B_d}{V_{Earth} B_m} \simeq 0.016 \quad \text{at LNGS latitude}$$

- Observed annual modulation amplitude in DAMA/LIBRA-phase1 in the (2–6) keV energy interval:  $(0.0097 \pm 0.0013)$  cpd/kg/keV
- Thus, the expected value of the diurnal modulation amplitude is  $\approx 1.5 \times 10^{-4}$  cpd/kg/keV.
- When fitting the *single-hit* residuals with a cosine function with period fixed at 24 h and phase at 14 h: all the diurnal modulation amplitudes  $A_d$  are compatible with zero at the present level of sensitivity.

$$A_d(2-6 \text{ keV}) < 1.2 \times 10^{-3} \text{ cpd/kg/keV (90\%CL)}$$

Model-independent result on possible diurnal effect in DAMA/LIBRA-phase1



Present experimental sensitivity is not yet enough for the expected diurnal modulation amplitude derived from the DAMA/LIBRA-phase1 observed effect.

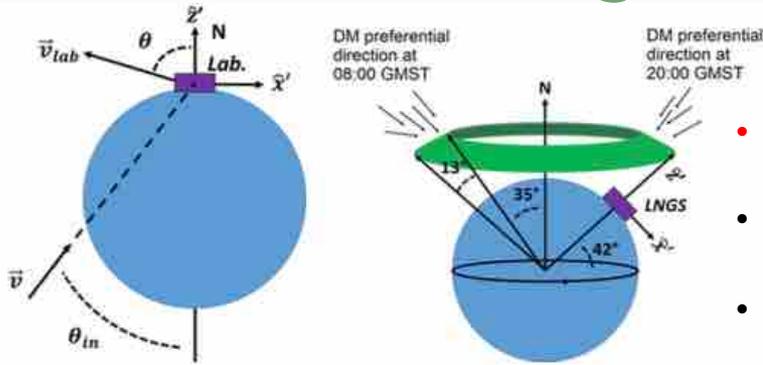
larger exposure DAMA/LIBRA-phase2 (+lower energy threshold) offers increased sensitivity to such an effect

## *Other signatures?*

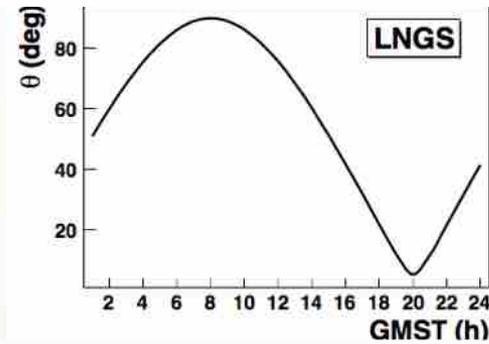
- *Diurnal effects*
- *Second order effects*
- *Shadow effects*
- *Directionality*
- *...*

# Earth shadowing effect with DAMA/LIBRA-phase1

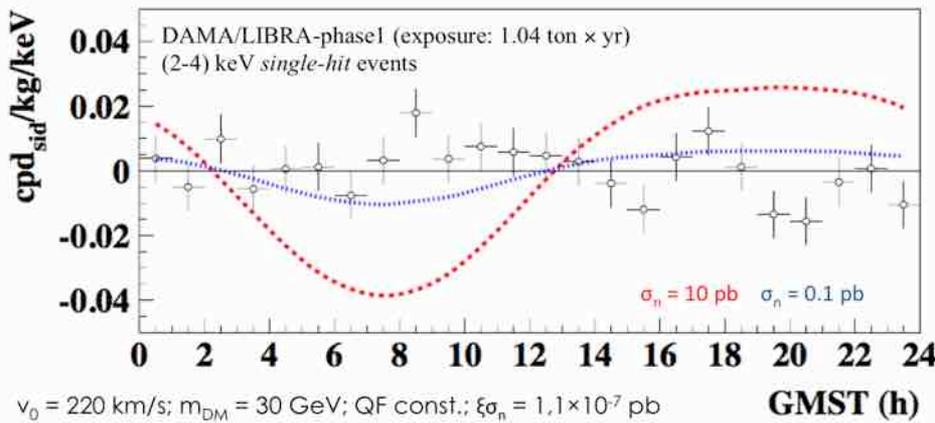
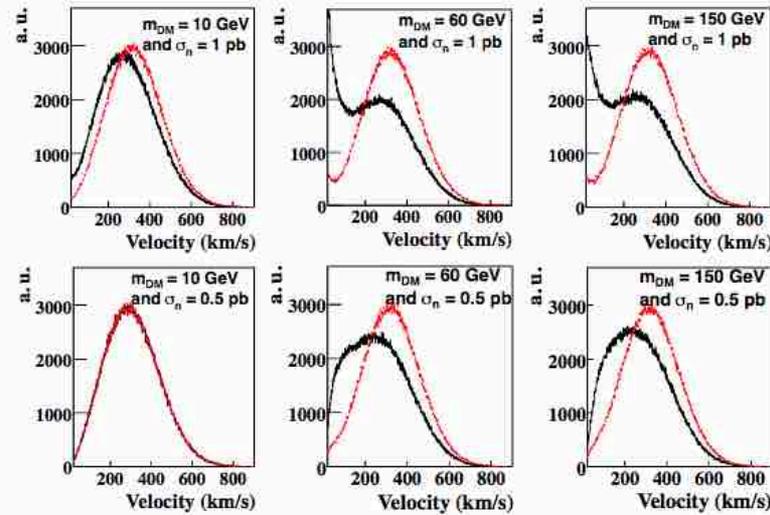
EPJC75(2015)239



- **Earth Shadow Effect** could be expected for DM candidate particles inducing nuclear recoils
- can be pointed out only for candidates with high cross-section with ordinary matter (low DM local density)
- would be induced by the variation during the day of the Earth thickness crossed by the DM particle in order to reach the experimental set-up



- DM particles crossing Earth lose their energy
- DM velocity distribution observed in the laboratory frame is modified as function of time (**GMST 8:00 black**; **GMST 20:00 red**)

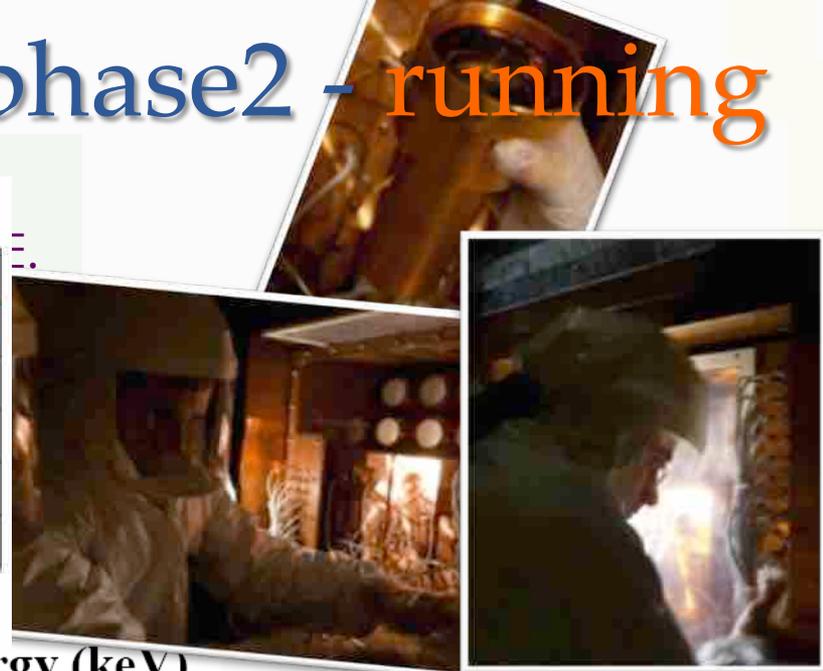
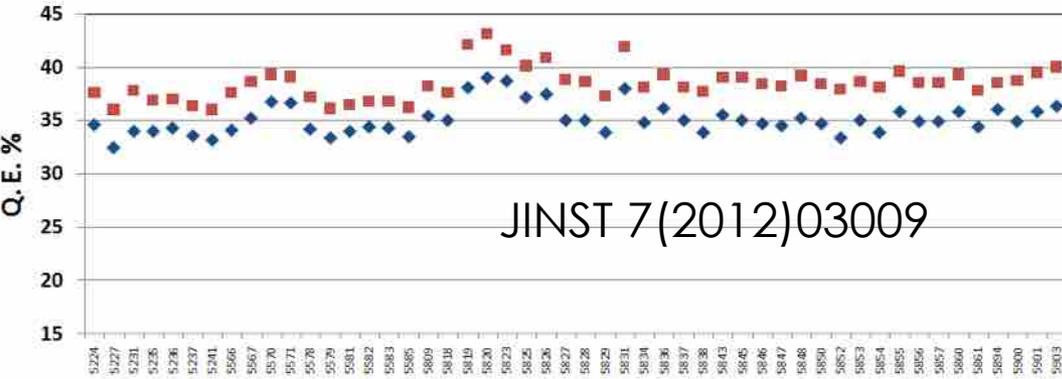


Taking into account the DAMA/LIBRA DM annual modulation result, allowed regions in the  $\xi$  vs  $\sigma_n$  plane for each  $m_{DM}$ .

# DAMA/LIBRA phase2 - running

## Quantum Efficiency features

■ Q.E. @ peak (%)    ◆ Q.E. @ 420 nm (%)



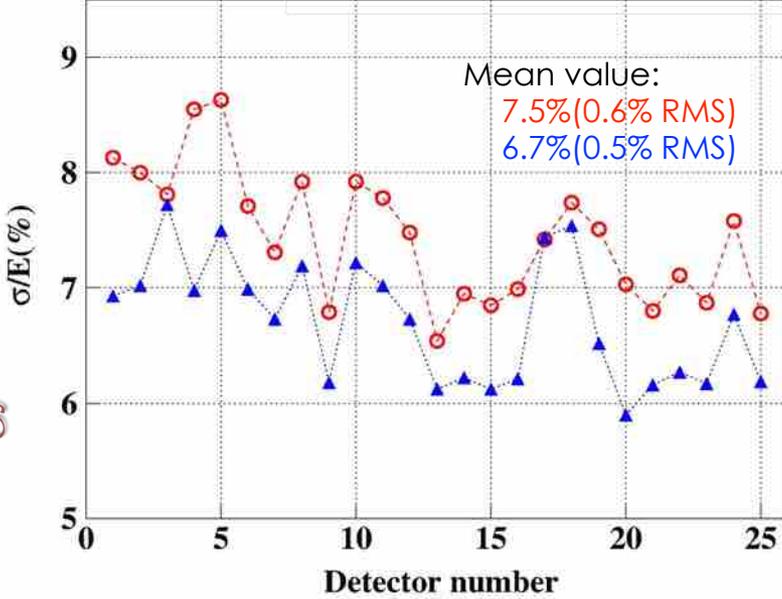
Residual Contamination

Serial number  
The limits are at 90% C.L.

Energy (keV)

PMT	Time (s)	Mass (kg)	<sup>226</sup> Ra (Bq/kg)	<sup>210</sup> Pb (Bq/kg)	<sup>235</sup> U (mBq/kg)	<sup>228</sup> Ra (Bq/kg)	<sup>232</sup> Th (mBq/kg)	<sup>40</sup> K (Bq/kg)	<sup>137</sup> Cs (mBq/kg)	<sup>60</sup> Co (mBq/kg)
Average			0.43	-	47	0.12	83	0.54	-	-
Standard deviation			0.06	-	16	0.02	17	0.16	-	-

Energy resolution



$\sigma/E$  @ 59.5 keV for each detector with new PMTs with higher quantum efficiency (blue points) and with previous PMT EMI-Electron Tube (red points).

### The light responses

Previous PMTs: 5.5-7.5 ph.e./keV  
New PMTs: up to 10 ph.e./keV

- To study the nature of the particles and features of related astrophysical, nuclear and particle physics aspects, and to investigate second order effects
- Special data taking for *other rare processes*

# DAMA/LIBRA phase 2 – data taking

- ✓ Fall 2012: new preamplifiers installed + special trigger modules.
- ✓ Calibrations 5 a.c.:  $\approx 1.03 \times 10^8$  events from sources
- ✓ Acceptance window eff. 5 a.c.:  $\approx 7 \times 10^7$  events ( $\approx 2.8 \times 10^6$  events/keV)

Annual Cycles	Period	Mass (kg)	Exposure	( $\alpha$ - $\beta^2$ )
I	Dec 2010 – Sept. 2011		commissioning	
II	Nov. 2, 2011 – Sept. 11, 2012	242.5	62917	0.519
III	Oct. 8, 2012 – Sept. 2, 2013	242.5	60586	0.534
IV	Sept. 8, 2013 – Sept. 1, 2014	242.5	73792	0.479
V	Sept. 1, 2014 – Sept. 9, 2015	242.5	71180	0.486
VI	Sept. 10, 2015 – Sept. 6, 2016	242.5	$\approx 70000$ (under analysis)	
VII	Sept 2016 –	242.5	running	

**PRELIMINARY**

Exposure expected for the first data release of DAMA/LIBRA-phase2:  $\approx 1 \text{ ton x yr}$

# Features of the DM signal

The importance of studying **second order effects** and the **annual modulation phase**

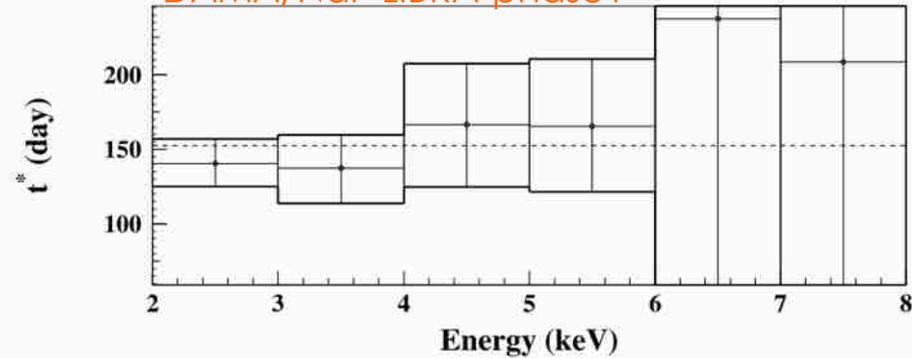
High exposure and lower energy threshold can allow further investigation on:

- the nature of the DM candidates
- possible diurnal effects on the sidereal time
- astrophysical models

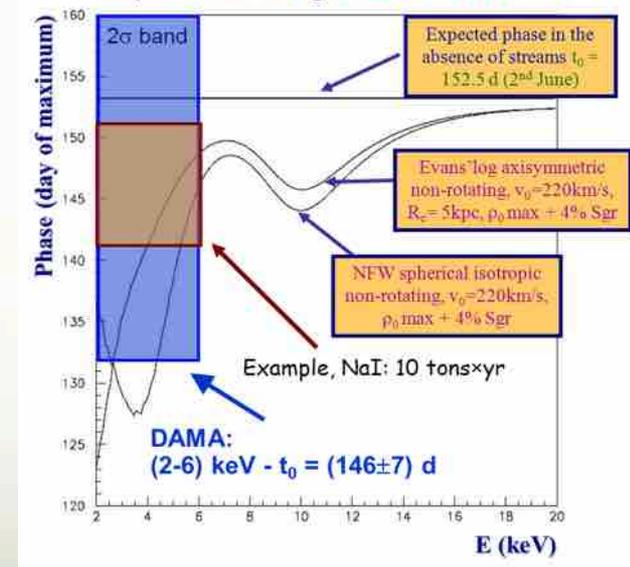
The annual modulation phase depends on :

- Presence of **streams** (as SagDEG and Canis Major) in the Galaxy
- Presence of **caustics**
- Effects of gravitational **focusing of the Sun**

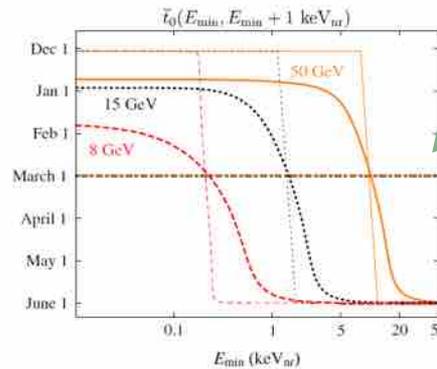
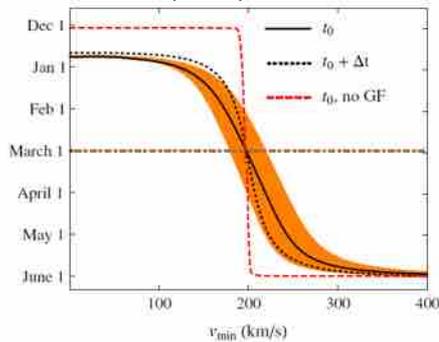
DAMA/NaI+LIBRA-phase1



The effect of the streams on the phase depends on the galactic halo model



PRL112(2014)011301



A step towards such investigations:

**DAMA/LIBRA-phase2** running with lower energy threshold

+ further possible improvements (DAMA/LIBRA-phase3) and **DAMA/1ton**

# Possible DAMA/LIBRA-phase3

- The light collection of the detectors can further be improved
- Light yields and the energy thresholds will improve accordingly

The strong interest in the low energy range suggests the possibility of a new development of **high Q.E. PMTs** with **increased radiopurity** to directly couple them to the DAMA/LIBRA crystals, **removing** the special radio-pure quartz (Suprasil B) light guides (10 cm long), which act also as optical window.

The presently-reached PMTs features, but not for the same PMT mod.:

- Q.E. around 35-40% @ 420 nm (NaI(Tl) light)
- radiopurity at level of 5 mBq/PMT ( $^{40}\text{K}$ ), 3-4 mBq/PMT ( $^{232}\text{Th}$ ), 3-4 mBq/PMT ( $^{238}\text{U}$ ), 1 mBq/PMT ( $^{226}\text{Ra}$ ), 2 mBq/PMT ( $^{60}\text{Co}$ ).

R&D efforts to obtain PMTs matching the best performances... **feasible**

No longer need for light guides (a 30-40% improvement in the light collection is expected)



## *Other signatures?*

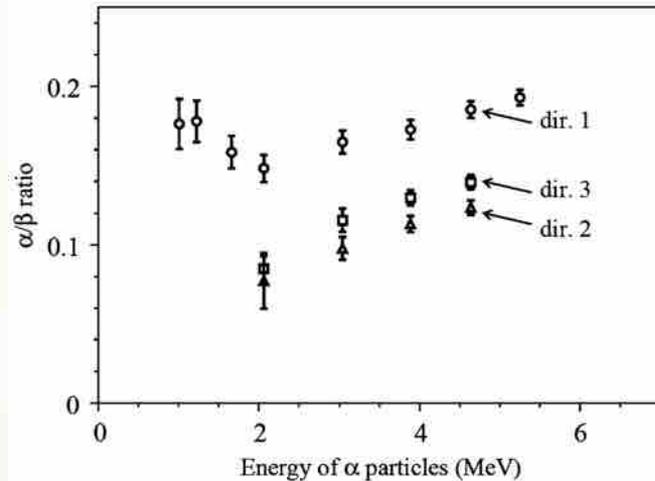
- *Diurnal effects*
- *Second order effects*
- *Shadow effects*
- *Directionality*
- *...*

# Directionality technique with crystals

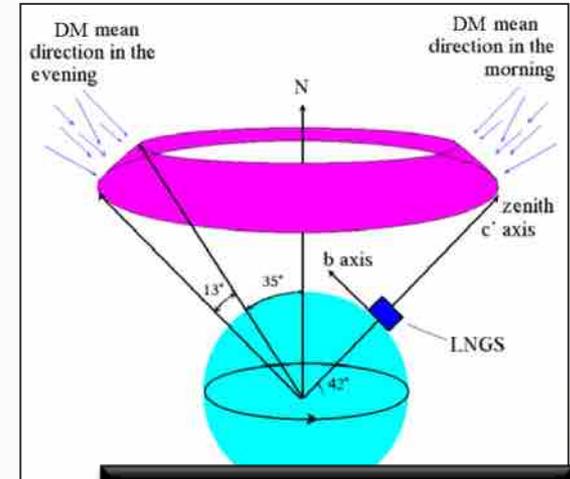
N. Cim. C15(1992)475, EPJC28(2003)203, EPJC73(2013)2276

- Only for candidates inducing just recoils
- Identification of the Dark Matter particles by exploiting the non-isotropic recoil distribution correlated to the Earth velocity

The ADAMO project: Study of the directionality approach with  $\text{ZnWO}_4$  anisotropic detectors



Nuclear recoils are expected to be strongly correlated with the DM impinging direction. This effect can be pointed out through the study of the variation in the response of anisotropic scintillation detectors during sidereal day.



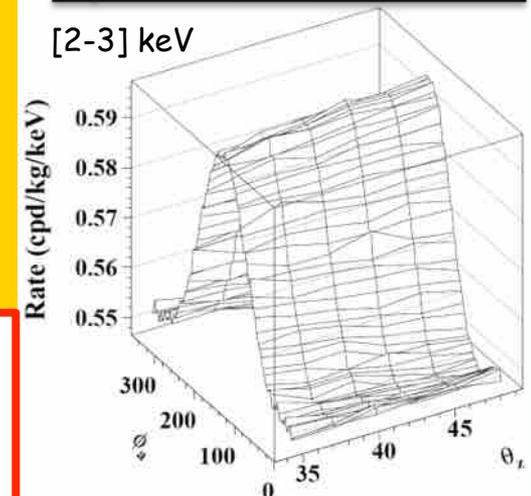
$$\sigma_p = 5 \times 10^{-5} \text{ pb}, m_{\text{DM}} = 50 \text{ GeV}$$

The light output and the pulse shape of  $\text{ZnWO}_4$  detectors depend on the direction of the impinging particles with respect to the crystal axes.

Both these anisotropic features can provide two independent ways to exploit the directionality approach.

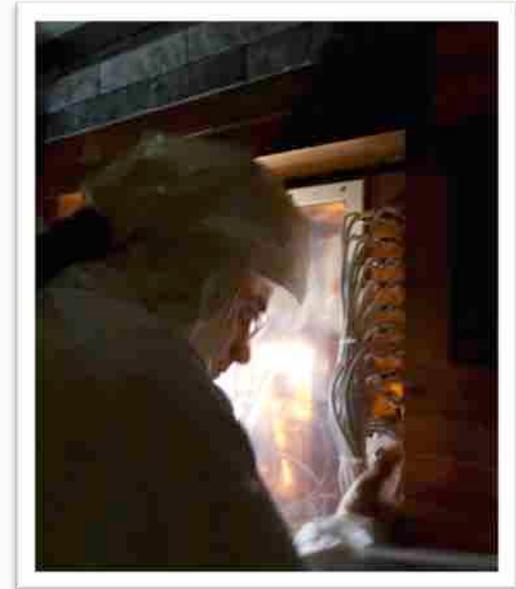
These and other competitive characteristics of  $\text{ZnWO}_4$  detectors could permit to reach sensitivity comparable with that of the DAMA/LIBRA positive result.

Presently running at ENEA-Casaccia with neutron generator to measure anisotropy at keV range.



# Conclusions

- Positive evidence for the presence of DM particles in the galactic halo supported at  **$9.3\sigma$**  C.L. (14 annual cycles DAMA/NaI and DAMA/LIBRA-phase1: 1.33 ton  $\times$  yr)
- Modulation parameters determined with high precision
- New investigation on different peculiarities of the DM signal exploited (**Diurnal Modulation** and **Earth Shadow Effect**)
- New corollary analysis on **Mirror Dark Matter**
- Full sensitivity to many kinds of DM candidates and interactions types (both inducing recoils and/or e.m. radiation), full sensitivity to low and high mass candidates



- **DAMA/LIBRA – phase2 in data taking** at lower software energy threshold (below 2 keV)
- Continuing investigations of rare processes other than DM
- **DAMA/LIBRA – phase3 R&D in progress**
- R&D for a possible DAMA/1ton set-up, proposed by DAMA since 1996, **continuing**
- Study of  $\text{ZnWO}_4$  scintillator for exploiting directionality technique **in progress**

